

# Northumbria Research Link

Citation: Kuridze, David, Mathioudakis, Mihalis, Jess, David, Shelyag, Sergiy, Christian, Damian, Keenan, Francis and Balasubramaniam, Karatholuvu (2011) Small-scale H $\alpha$  jets in the solar chromosphere. *Astronomy & Astrophysics*, 533. A76. ISSN 0004-6361

Published by: EDP Sciences

URL: <http://dx.doi.org/10.1051/0004-6361/201117427> <<http://dx.doi.org/10.1051/0004-6361/201117427>>

This version was downloaded from Northumbria Research Link:  
<http://nrl.northumbria.ac.uk/26903/>

Northumbria University has developed Northumbria Research Link (NRL) to enable users to access the University's research output. Copyright © and moral rights for items on NRL are retained by the individual author(s) and/or other copyright owners. Single copies of full items can be reproduced, displayed or performed, and given to third parties in any format or medium for personal research or study, educational, or not-for-profit purposes without prior permission or charge, provided the authors, title and full bibliographic details are given, as well as a hyperlink and/or URL to the original metadata page. The content must not be changed in any way. Full items must not be sold commercially in any format or medium without formal permission of the copyright holder. The full policy is available online: <http://nrl.northumbria.ac.uk/policies.html>

This document may differ from the final, published version of the research and has been made available online in accordance with publisher policies. To read and/or cite from the published version of the research, please visit the publisher's website (a subscription may be required.)

[www.northumbria.ac.uk/nrl](http://www.northumbria.ac.uk/nrl)



# Small-scale H $\alpha$ Jets in the Solar Chromosphere

D.Kuridze<sup>1</sup>, M.Mathioudakis<sup>1</sup>, D.B.Jess<sup>1</sup>, S.Shelyag<sup>1</sup>, D.J.Christian<sup>2</sup>, F.P.Keenan<sup>1</sup>, K.S.Balasubramaniam<sup>3</sup>

<sup>1</sup> Astrophysics Research Centre, School of Mathematics and Physics, Queen's University Belfast, Belfast, BT7 1NN, Northern Ireland, U.K.

<sup>2</sup> Department of Physics and Astronomy, California State University, Northridge, CA 91330, USA

<sup>3</sup> Air Force Research Laboratory, Solar and Solar Disturbances, Sunspot, NM 88349, USA

received / accepted

## ABSTRACT

*Aims.* High temporal and spatial resolution observations from the Rapid Oscillations in the Solar Atmosphere (ROSA) multiwavelength imager on the Dunn Solar Telescope are used to study the velocities of small-scale H $\alpha$  jets in an emerging solar active region.

*Methods.* The dataset comprises of simultaneous imaging in the H $\alpha$  core, Ca II K, and G band, together with photospheric line-of-sight magnetograms. Time-distance techniques are employed to determine projected plane-of-sky velocities.

*Results.* The H $\alpha$  images are highly dynamic in nature, with estimated jet velocities as high as 45 km s<sup>-1</sup>. These jets are one-directional, with their origin seemingly linked to underlying Ca II K brightenings and G-band magnetic bright points.

*Conclusions.* It is suggested that the siphon flow model of cool coronal loops is suitable for the interpretation of our observations. The jets are associated with small-scale explosive events, and may provide a mass outflow from the photosphere to the corona.

**Key words.** Sun: activity – Sun: chromosphere – Sun: faculae, plages – Sun: photosphere – Sun: surface magnetism

## 1. Introduction

The solar chromosphere is permeated by a wide range of highly dynamic features including jets, fibrils, mottles, spicules, Ellerman bombs and H $\alpha$  surges. In particular, chromospheric jets are one of the most important, yet also most poorly understood phenomena of the Sun's magnetic atmosphere. Observed chromospheric jet velocities are generally measured to have values of around  $\sim 20\text{--}40$  km s<sup>-1</sup> (Tsiropoula & Tziotziou 2004; Chae et al. 2000; Lin et al. 2008). Apart from these typical speeds, high velocity flows ( $>100$  km s<sup>-1</sup>) are also observed in some chromospheric surges and blobs (Foukal 2004; van Noort et al. 2006).

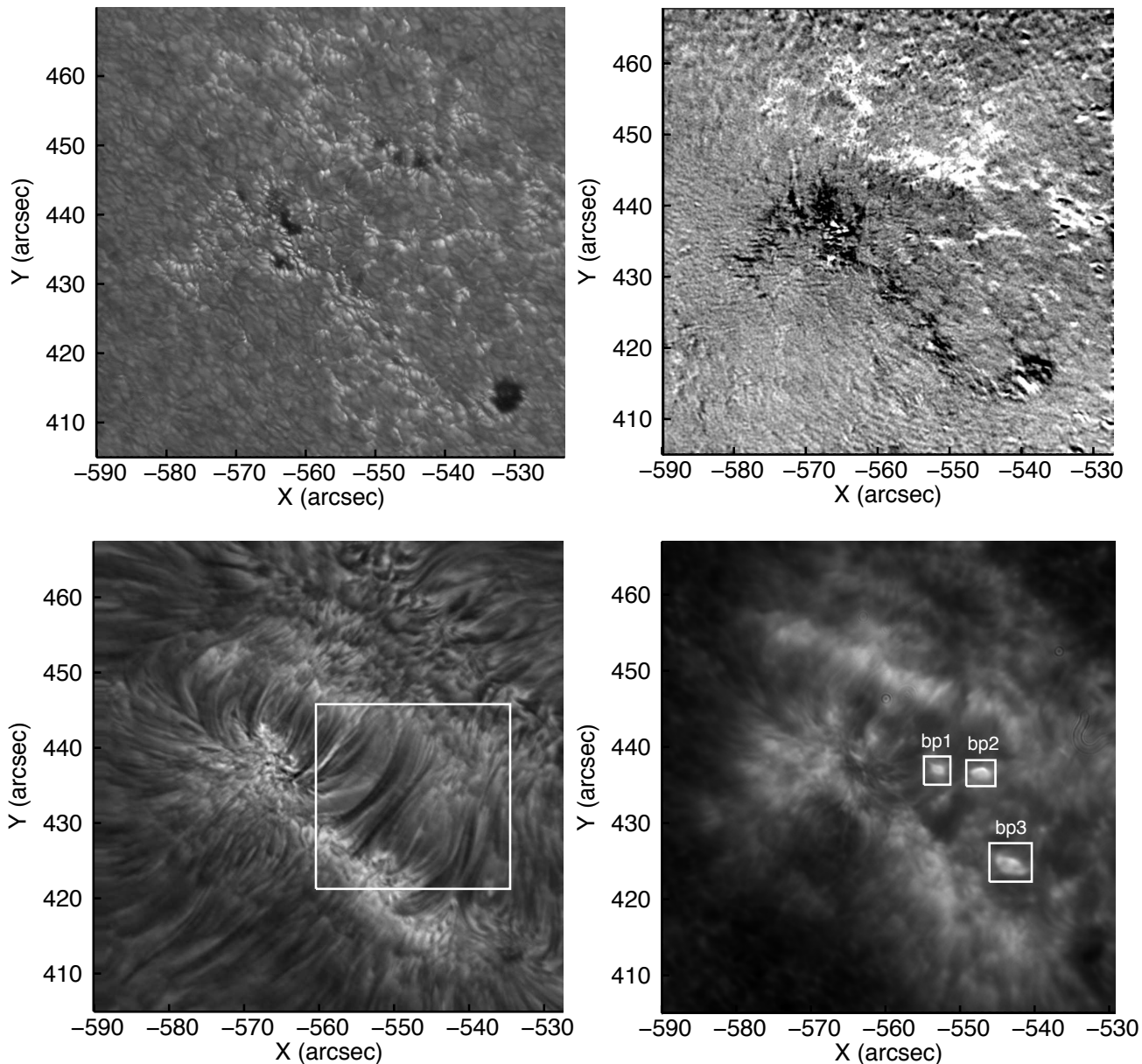
Several previous studies indicate that the energy required to drive these chromospheric jets can be released by magnetic reconnection occurring low in the solar atmosphere (Canfield et al. 1996; Chae et al. 2003; Chae 2007; Shibata et al. 2007). Yoshimura et al. (2003) found a close correlation between H $\alpha$  surges, Transition Region and Coronal Explorer (TRACE) brightenings in the 1600 Å waveband, and magnetic flux cancellations surrounding ephemeral regions. More recently, Morita et al. (2010) observed chromospheric anemone jets in Ca II K<sub>1</sub> and K<sub>2</sub> lines as intensity peaks in the light curves. They also found cancellation of magnetic flux near the jet area. These results strongly suggest that chromospheric jets may be generated by magnetic reconnection, similar to that occurring on a larger scale in the corona. Also, the multi-wavelength studies of Yoshimura et al. (2003) and Brooks et al. (2007) indicate that chromospheric surges are indeed driven by magnetic reconnection. Alternatively, Hansteen et al. (2006) suggest that the formation of jets, linked to dynamic fibrils,

mottles and spicules, could be driven by magneto-acoustic shocks that leak upward through the chromosphere as a result of convective flows in the photosphere and global p-mode oscillations.

Recently, a multi-wavelength study involving Solar Dynamic Observatory (SDO) and Hinode observations has highlighted the importance of chromospheric jets and spicules for providing the mass supply for the corona (De Pontieu et al. 2011). However, spectroscopic observations can only provide the line-of-sight component of velocity, with any small-scale Doppler phenomena cancelling out if the spatial resolution is not sufficiently high. To investigate the small-scale structure, here we present high spatial and temporal resolution observations of an emerging active region in H $\alpha$  core, Ca II K, G band and line-of-sight magnetograms. In the following sections, we describe the details of the observations and the image processing (§2), the analysis and results (§3), and finally discuss and summarise possible interpretations of our findings (§4).

## 2. Observations & Data Reduction

The observations were obtained between approximately 15:00–16:00 UT on 2009 May 31 with the Rapid Oscillations in the Solar Atmosphere (ROSA; Jess et al. 2010b) imaging system, mounted on the Dunn Solar Telescope (DST) at the National Solar Observatory, New Mexico, USA. High-order adaptive optics were used throughout the observations. The photospheric target under investigation comprises of a multitude of magnetic pores and associated faculae, located at heliocentric coordinates ( $-540''$ ,  $435''$ ). Our dataset in-

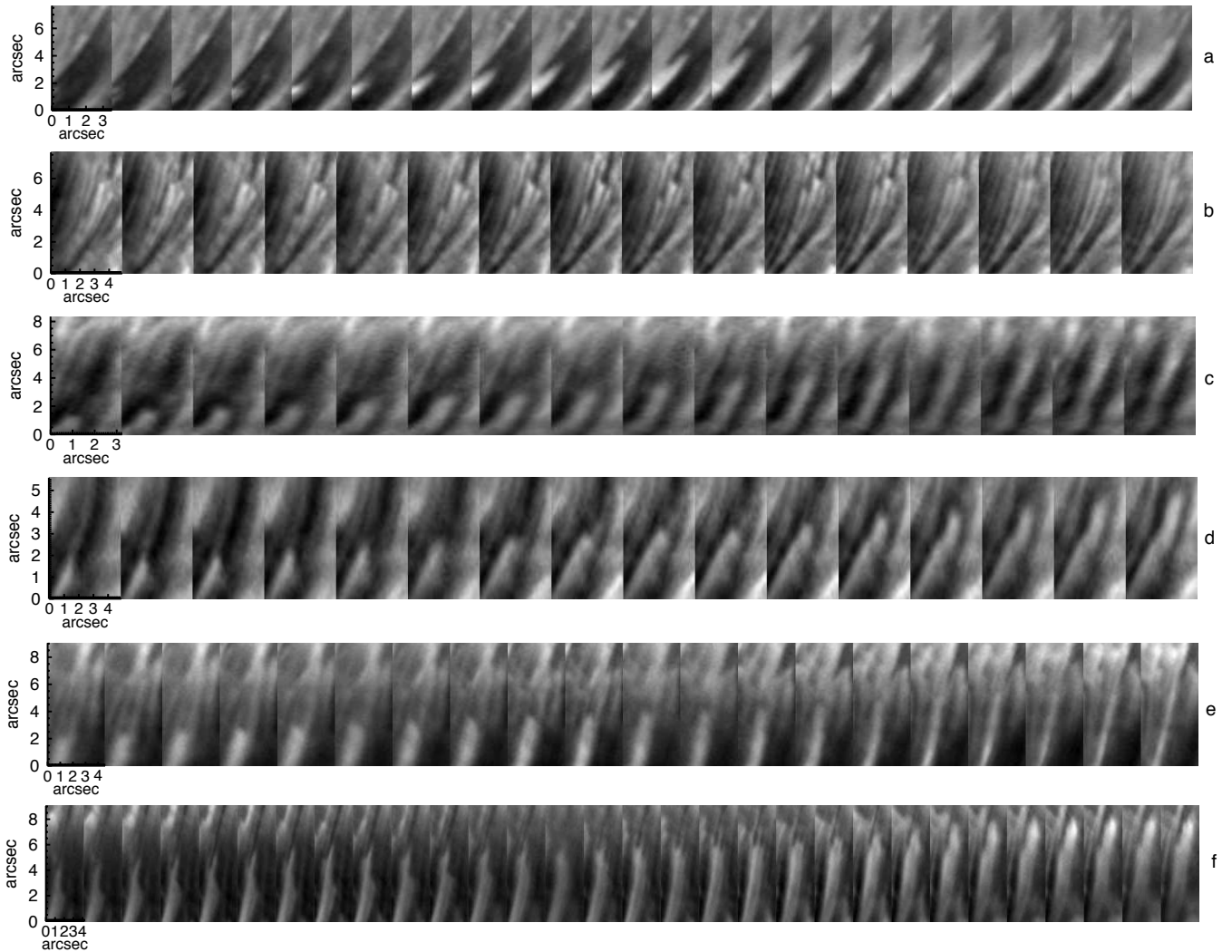


**Fig. 1.** Simultaneous ROSA images of the G band (top left), line-of-sight magnetogram (top right), H $\alpha$  core (bottom left), and Ca II K core (bottom right). The magnetogram colour scale is in Gauss, where white indicates positive magnetic polarities, while black demonstrates polarities which are negative. Artificial saturation at  $\pm 400$  G is implemented to highlight regions of weak magnetic polarity on the solar surface. The white rectangle in the H $\alpha$  core image indicates the approximate position where most of the jets are observed. White rectangles in the Ca II K image, identified by ‘bp1’, ‘bp2’, and ‘bp3’, are the bright points selected for temporal analysis. Axes are in heliocentric arcseconds, where  $1'' \approx 725$  km.

cludes simultaneous imaging in the G band, H $\alpha$  core, Ca II K and line-of-sight (LOS) magnetograms. A spatial sampling of  $0.069''/\text{pixel}$  was selected, matching the diffraction limit of the DST to that of the G band, resulting in a field-of-view of approximately  $69'' \times 69''$ .

The images obtained were processed with speckle reconstruction algorithms (Wöger et al. 2008), with the removal of large-scale seeing distortions achieved by destretching the data relative to simultaneous high-contrast continuum images. The algorithms utilised 32 images per reconstruction, resulting in an effective cadence of 8.4 s for H $\alpha$ , Ca II K, and LOS magnetograms, while the reconstructed G-band cadence was 1.1 s. Magnetograms were constructed

as normalised to their sum difference images of left- and right-hand circularly polarized light obtained 125 mÅ into the blue wing of the magnetically-sensitive Fe I absorption line at 6302.5 Å. A blue-wing offset was required to minimise granulation contrast, while conversion of the filtergram into units of Gauss was performed using simultaneous SoHO/MDI magnetograms (see discussion in Jess et al. 2010a). ROSA images in G band, H $\alpha$  core, Ca II K and LOS magnetograms are shown in Fig. 1.



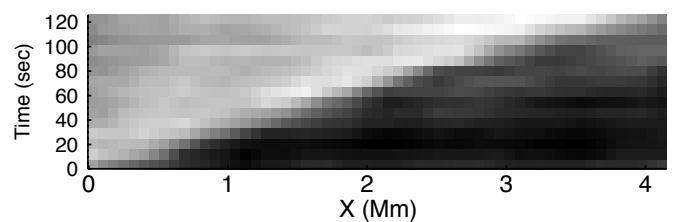
**Fig. 2.** Six examples of H $\alpha$  jets observed within the white rectangle in the lower-left panel of Fig. 1. Each successive image is separated by a time of 8.4 s, while the length of the jet structures range between approximately 3000–6000 km.

### 3. Analysis and discussion

The area under investigation (see boxed region in the lower-left panel of Fig. 1) consists of a multitude of highly dynamic thread-like structures, which connect areas of opposite magnetic polarity, forming a “chromospheric arcade”.

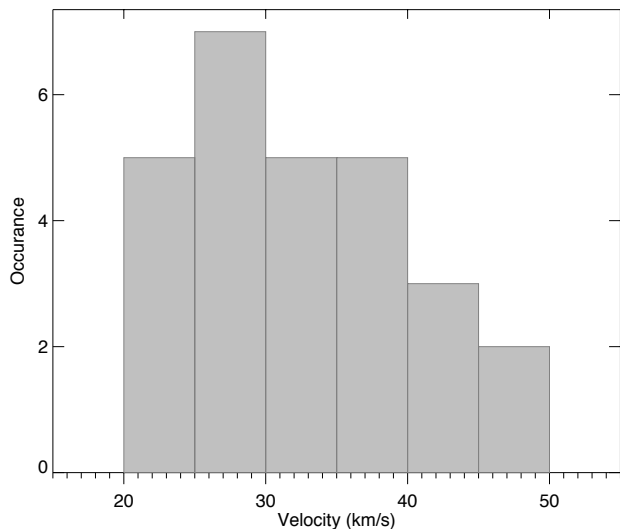
A thorough examination of the H $\alpha$  core images reveals a total of 27 jet-like structures, some of which are displayed in Fig. 2. The velocities of the jets are determined using space-time slices ( $x-t$  plots), generated along the path of the jet trajectories (Fig. 3). Velocity estimates for these jets are in the range 20–45 km s $^{-1}$ , with a histogram of their occurrence plotted in Fig. 4. Detailed properties for the six most prominent jets (see Fig. 2) are listed in Table 1.

Almost half of the H $\alpha$  jets originate from locations co-spatial with Ca II K and G-band bright points. Three examples of bright points are outlined in the lower-right panel of Fig. 1 using white rectangles and labelled as ‘bp1’, ‘bp2’, and ‘bp3’. All six jets shown in Fig. 2, as well six additional jets, are co-spatial with these bright features, and the evolution of the Ca II K bright point intensities indicates a possible correlation with the H $\alpha$  jets. Ca II K intensity enhancements and strong isolated peaks are visible in



**Fig. 3.** H $\alpha$  time-distance diagram of jet *a*, taken from Fig. 2. The gradient of the diagonal ridge in this  $x-t$  slice corresponds to the jet velocity. The velocity of this jet is 34 km s $^{-1}$ .

the bright point light curves shown in Fig. 5. In the light curve bp1 (top panel), the strongest peak around 15:20 UT is followed by three H $\alpha$  jets occurring at  $\sim$  15:21, 15:24, 15:26 UT (crosses). From 15:29, there is no significant enhancement in the Ca II K intensity and correspondingly no H $\alpha$  events. There is an enhanced intensity for the first few minutes in the light curve of bp2. During this time interval,



**Fig. 4.** Velocity distribution of the 27 H $\alpha$  jets.

**Table 1.** Detailed parameters of the jets displayed in Fig. 2.

Event name	Starting time (UT)	Duration of the event (sec)	Velocity (km/s)
Jet a	15:24:46	160	34
Jet b	15:03:53	142	20
Jet c	15:31:36	135	35
Jet d	15:33:08	145	38
Jet e	15:18:20	177	21
Jet f	15:05:27	250	23

four strong H $\alpha$  jets are generated at the location of bp2. Furthermore, one of the most isolated intensity peaks in the light curve of bp2 is around 15:25, and there is an H $\alpha$  jet observed very close to that time. Two other jets associated with the bp2 occur between 15:31 and 15:33. A sharp decrease in the intensity of bp2 occurs after this time. In the light curve of bp3, the peaks are located at around 15:10 and 15:18 UT. The eruption of two jets (e,f in the Fig. 2) are observed near these peaks. H $\alpha$  jets appear to initiate at times when the Ca II K intensity rises, suggesting a possible link between the two phenomena. It must be noted that there are some bright peaks in the light curves of considered bright points (Fig. 5) that are not accompanied by H $\alpha$  jet events. In these cases the jets may not be visible due to overlapping H $\alpha$  structures. The remaining jets, of the 27 analysed in this work, also correspond to Ca II K brightenings, however, these brightenings are not observed as localised, simple, fine areas, and makes them somewhat more difficult to disentangle.

Observed 27 jets are initiated in about 9 distinct, small-scale ( $\sim 5'' \times 5''$ ) regions. As we have already mentioned, almost half of them originate from the three bright points labelled in the bottom right panel in Fig. 1. This fact indicates repetitive nature of the observed jets, however, no well-defined period was found.

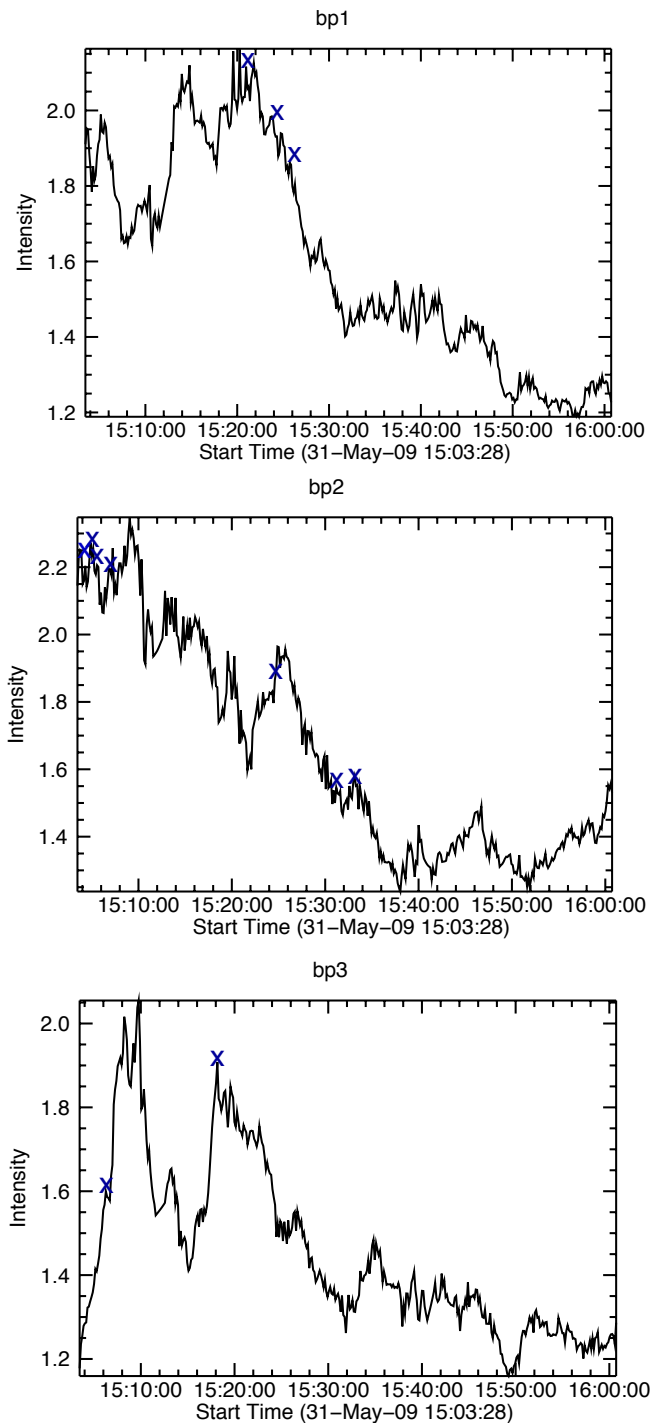
Light curves of background quiescent Ca II K regions do not indicate any intensity peaks or significant enhancements. We note that small amplitude intensity fluctuations consistent with the well-known 3-5 minute acoustic modes

are observed in both the quiescent and bright point regions (see e.g. Fleck & Schmitz 1991; Kariyappa et al. 2001). However, the intermittent nature of the jets and their supersonic velocities suggest that they are not associated with the acoustic mode oscillations.

Striking similarities between H $\alpha$  core observations and TRACE Fe IX/X images suggest that H $\alpha$  may be used as a proxy for coronal material. These similarities may be attributed to reduced H $\alpha$  opacity, excess emissivity, or a combination of both (Rutten 2007). A cool loop model may therefore be appropriate for chromospheric structures.

Doyle et al. (2006) suggested that one-directional plasma flow in a cold coronal loop ( $10^5$ – $10^6$  K) can be interpreted in terms of a short-lived siphon flow which may arise due to a non-linear heating pulse at one of the loop footpoints. The supersonic flow velocities are determined by the duration of the pulse, in addition to its associated energy input. All our presented H $\alpha$  jets are one-directional, and propagate from negative (bottom portion of the white rectangle in the lower-left panel of Fig. 1) to positive polarities (upper section of the white rectangle in the lower-left panel of Fig. 1). The established supersonic velocities are very similar to those predicted by the siphon flow model. Furthermore, other similarities between observations of Doyle et al. (2006) and our findings, such as the comparable time scales of the observed events and also morphological similarities of the observed structures motivated us to interpret our observation by means of siphon flow model. We note that Uitenbroek et al. (2006) used spectroheliograms of the Ca II IR triplet (8542.1 Å) to detect a siphon flow associated with a magnetic pore. Line of sight velocities as high as 27 km s $^{-1}$  were detected.

Using the heating rate and plasma flow velocity given by the coronal loop simulation of Doyle et al. (2006), we can estimate that the rate of energy conversion from the source to the kinetic energy of the flow in the loop is  $\sim 0.1$ . This estimate is based on an ambient coronal density of  $10^{-13}$  g  $\cdot$  cm $^{-3}$ . If we assume (i) that the same rate applies in the chromosphere, (ii) a chromospheric density of  $10^{-11}$  g  $\cdot$  cm $^{-3}$  and (iii) the chromospheric jet velocity of 40 km s $^{-1}$  as estimated from the observations, we obtain the total heating energy of about 10 erg  $\cdot$  cm $^{-3}$ . Assuming the duration of energy release in the chromospheric loop footpoint of the order of ten seconds, we obtain a heating rate  $h \sim 1$  erg  $\cdot$  cm $^{-3}$ s $^{-1}$ , a value significantly higher than the one estimated for the coronal loop. This is not surprising as more energy is required to propel the denser chromospheric jets to velocities comparable to those in the coronal model. The volume of the Ca II K bright points (bottom right panel of the Fig. 1), which are the footpoints of the H $\alpha$  jets, can be estimated as  $10^{24}$  cm $^3$  (assuming a size of approximately  $2'' \times 2''$  and a height of  $\sim 1''$ ). Thus, the estimated released energy for the heating rate calculated above is about  $10^{24}$  erg. The total mass of the jet ( $\sim 10^{12}$  g) is estimated from the chromospheric density and average volume of the jet. In order to bring the total mass up to a height of about  $1.5 \times 10^3$  km an energy of  $\sim 10^{24}$  erg is needed. These values are comparable to the estimated energy stored in the bright point volume. We emphasize that the above estimates are approximate and a detailed simulation is required to account for all processes involved in this energy release.



**Fig. 5.** Light curves of the Ca II K bright points outlined with rectangles in the lower-right panel of Fig. 1. ‘x’ symbols indicate the time when a jet event originated from that particular bright point.

#### 4. Concluding Remarks

We present evidence for small-scale, one directional H $\alpha$  jets in the solar chromosphere. Jet velocities as high as 45 km s $^{-1}$  are found. Observation shows the relations between H $\alpha$  jets and Ca II K bright points. We interpret this phenomena by means of a coronal loop siphon flow model, whereby one of the loop footpoints is driven by a localised, non-linear

heating pulse (Doyle et al. 2006). The energy required to drive jets is estimated as  $10^{24}$  erg.

The jets presented here appear to be associated with small-scale explosive events, and may provide a mechanism to support mass outflow from the chromosphere out into the corona.

*Acknowledgements.* Observations were obtained at the National Solar Observatory, operated by the Association of Universities for Research in Astronomy, Inc (AURA) under agreement with the National Science Foundation. This work is supported by the Science and Technology Facilities Council (STFC), with DBJ particularly grateful for the award of an STFC post-doctoral fellowship. We thank the Air Force Office of Scientific Research, Air Force Material Command, USAF for sponsorship under grant number FA8655-09-13085.

#### References

- Brooks, D. E., Kurokawa, H., & Berger, T. E. 2007, *ApJ*, 656, 1197  
 Canfield, R. C., Reardon, K. P., Leka, K. D., Shibata, K., Yokoyama, T., & Shimojo, M. 1996, *ApJ*, 464, 1016  
 Chae, J., Denker, C., Spirock, T., Wang, H., & Goode, P. 2000, *Sol. Phys.*, 195, 333  
 Chae, J., Moon, Y.-J., & Park, S.-Y. 2003, *JKAS*, 36, S13  
 Chae, J. 2007, in *ASP Conf. Ser. 369, New Solar Physics with Solar-B Mission*, ed. K. Shibata, S. Nagata, & T. Sakurai (San Francisco, CA: ASP), 243  
 De Pontieu, B., McIntosh, S.W., Carlsson, M., Tarbell, T.D., Boerner, P., Martinez-Sykora, J., Schrijver, C.J., Title, A.M., 2011, *Science*, 331, 55  
 Doyle, J. G., Taroyan, Y., Ishak, B., Madjarska, M. S., & Bradsgaw, 2006, *A&A*, 452,1075  
 Fleck, B., & Schmitz, F. 1991, *A&A*, 250, 235  
 Foukal, P. 2004, *Solar Astrophysics* (2nd ed; Weinheim: Wiley-VCH)  
 Hansteen, V. H., De Pontieu, B., Rouppe van der Voort, L., van Noort, M., & Carlsson, M. 2006, *ApJ*, 647, L73  
 Jess, D. B., Mathioudakis, M., Christian, D. J., Crockett, P. J., & Keenan, F. P. 2010a, *ApJ*, 719, L134  
 Jess, D. B., Mathioudakis, M., Christian, D. J., Keenan, F. P., Ryans, R. S. I., & Crockett, P. J. 2010b, *Sol. Phys.*, 261, 363  
 Kariyappa, R., Varghese, B. A., & Curdt, W. 2001, *A&A*, 374, 691  
 Lin, Y., Martin, S. F., Engvold, O., Rouppe van der Voort, L. H. M., & van Noort, M. 2008, *Adv. Space Res.*, 42, 803  
 Morita, S., Shibata, K., Ueno, S., Ichimoto, K., Kitai, R., & Otsuji, K. 2010, *PASJ*, 62, 901  
 Rutten, R.J., 2007, *ASP Conference Series*, eds. P.Heinzel, I. Dorotovic, & R.J. Rutten, San Francisco, p27  
 Shibata, K., et al. 2007, *Science*, 318, 1591  
 Tsiropoula G., & Tziotziou K. 2004, *A&A*, 424, 279  
 Uitenbroek, H., Balasubramaniam, K.S., Tritschler, A., *ApJ*, 645, 776  
 van Noort, M., & Rouppe van der Voort, L. 2006, *ApJ*, 648, L67  
 Wöger, F., von der Lühe, O., & Reardon, K., 2008, *A&A*, 488, 375  
 Yoshimura, K., Kurokawa, H., Shimojo, M., & Shine, R. 2003, *PASJ*, 55, 313